



Early-type galaxies in the Antlia cluster: global properties

Juan P. Calderón^{1,2,3}★, Lilia P. Bassino,^{1,2,3} Sergio A. Cellone,^{1,3,4} Matías Gómez⁵ and Juan P. Caso^{1,2,3}

¹Consejo Nacional de Investigaciones Científicas y Técnicas, Godoy Cruz 2290, C1425FQB, Ciudad Autónoma de Buenos Aires, Argentina

²Instituto de Astrofísica de La Plata (CCT La Plata – CONICET – UNLP), Paseo del Bosque S/N, B1900FWA La Plata, Argentina

³Facultad de Ciencias Astronómicas y Geofísicas de la Universidad Nacional de La Plata, Paseo del Bosque S/N, B1900FWA La Plata, Argentina

⁴Complejo Astronómico El Leoncito (CONICET – UNLP – UNC – UNSJ), San Juan, Argentina

⁵Departamento de Ciencias Físicas, Facultad de Ciencias Exactas, Universidad Andres Bello, Santiago, Chile

Accepted 2020 July 8. Received 2020 July 8; in original form 2020 April 22

ABSTRACT

We present an extension of our previous research on the early-type galaxy population of the Antlia cluster ($d \sim 35$ Mpc), achieving a total coverage of $\sim 2.6 \text{ deg}^2$ and performing surface photometry for ~ 300 galaxies, 130 of which are new uncatalogued ones. Such new galaxies mainly fall in the low surface brightness (LSB) regime, but there are also some lenticulars (S0), which support the existence of unique functions that connect bright and dwarf galaxies in the scaling relations. We analyse the projected spatial distribution of galaxies up to a distance of ~ 800 kpc from NGC 3268, the adopted centre, as well as the radial velocity distribution and the correlation between galaxy colour and effective radius with the projected spatial distribution. We also obtain the luminosity function of the early-type galaxies and the distribution of stellar masses using the T_1 -band magnitudes and adopted mass-luminosity ratios. Additionally, we correlate the central galaxy distribution with an X-ray emission map from the literature. Based on the analysis of the radial velocities and galaxy colour distributions, we find that galaxies redder than the colour–magnitude relation (CMR) have a velocity distribution strongly concentrated towards the values of the dominant galaxies and are homogeneously distributed throughout the cluster. Those bluer than the CMR, in turn, have a much more extended radial velocity distribution and are concentrated towards the centre of the cluster. We also identify 12 candidates to ultra diffuse galaxies, which seem to be split into two families, and speculate about their origins in the context of the cluster structure.

Key words: galaxies: clusters: general – galaxies: clusters: individual: Antlia – galaxies: dwarf – galaxies: elliptical and lenticular, cD – galaxies: fundamental parameters.

1 INTRODUCTION

Over the last years, low-luminosity galaxies have received increasing attention from the community, due to newly discovered galaxies in nearby clusters, with unexpected surface brightness and size properties. The complexity attained by modern Λ CDM models in reference to baryonic effects (Tissera et al. 2010; Vera-Ciro et al. 2011; Gargiulo et al. 2019) and the new observational evidence brought by larger detector arrays – MOSAIC II,¹ Dark Energy Camera (DECam Flaugher et al. 2015) and OmegaCAM Camera (OmegaCAM Kuijken et al. 2002) – have led the research on the distribution and properties of these newly detected galaxies. These galaxies with extreme properties represent new challenges for the theoretical models of galaxy formation and evolution.

While early-type galaxies (ETGs) in the lowest surface brightness regime have been usually classified as dwarf ellipticals (dEs), at fainter magnitudes ($M_V \gtrsim -16$ mag), ETGs are represented by dwarf spheroidals (dSphs, Grebel, Gallagher & Harbeck 2003), which show effective radii (r_e) smaller than ~ 1 kpc, and are supposed to have experienced intense gas and metal removal due to galactic winds (Kirby, Martin & Finlator 2011). Still lower in the luminosity scale,

the ultrafaint dwarf galaxies (UFDGs, Willman et al. 2005) are characterized for being the most dark matter dominated, the most metal-poor, and least chemically evolved stellar systems ever investigated (Simon 2019). The stellar systems with the lowest surface brightness, observed in nearby clusters as Virgo, Fornax, and Coma, are the so-called ultra diffuse galaxies (UDGs, van Dokkum et al. 2015; Lim et al. 2018). They have large radii ($r_e \gtrsim 1$ kpc), moderate stellar masses, and their low surface brightnesses ($\mu_{0,V} \sim 27 \text{ mag arcsec}^{-2}$) make them accessible only to very deep observations (Eigenthaler et al. 2018). Amorisco & Loeb (2016) show that the origin of UDGs in clusters could be well explained by the classical model of disc formation in haloes with higher initial angular momentum.

Modern imaging surveys – e.g. Next Generation Virgo Cluster Survey (NGVS, Ferrarese et al. 2012), Fornax Deep Survey (FDS, Venhola et al. 2017), or Next Generation Fornax Survey (NGFS, Eigenthaler et al. 2018) – allow to perform very deep multiband photometry, covering at the same time a projected area on the sky similar to that attained by old photographic plates. This enables the study of the projected spatial distribution of the different populations that coexist in the clusters, even outside the virial radius. Despite that, at such low surface brightness (LSB) levels, cluster membership through spectroscopic redshift determinations is very difficult to achieve. Anyway, such surveys have provided us with important data to study the effects of tidal interactions in galaxy evolution. The most relevant environmental effect for the evolution of early-type dwarfs is

* E-mail: juan.pablo.calderon0@gmail.com

¹ <https://www.noao.edu/ctio/mosaic/>

harassment (Moore et al. 1996), which is related with repeated tidal encounters between the infalling galaxy and other cluster members or the cluster gravitational potential. This could produce a galaxy mass loss that depends on the orbit of the infalling galaxy within the cluster, and it is expected to be more efficient for low-mass galaxies (Smith, Davies & Nelson 2010). On the other hand, the motion of the infalling galaxy across the intracluster medium could produce the loss of the galaxy atomic gas and, consequently, the interruption of star formation (Gunn & Gott 1972). Hydrodynamical simulations (and references therein, Roediger & Brüggen 2008) indicate that this effect is not only due to gravitational interactions but also due to mixing turbulence and viscous effects. Additionally, an alternative effect, called *strangulation*, has been proposed, in which the infall of cool gas into the galaxy is stopped (Balogh, Navarro & Morris 2000). This can happen when a subhalo of dark matter is created from a previously existing larger halo. The infall of gas associated to the subhalo can be cut, which results in a gradual decrease in the stellar formation, causing the galaxy colour to become redder (Wang et al. 2009). The result of the strangulation effect is a galaxy with higher metallicity and stellar mass than its progenitor (Peng, Maiolino & Cochrane 2015). The efficiency of those physical processes is directly related with the parameters that define the structure and dynamics of each galaxy, such as total mass, relative orbits, and velocities, and also their spatial distribution inside the cluster potential. The so-called *morphology–density relation* (Dressler et al. 1997; van der Wel et al. 2010), which locates the ETGs mainly in the central areas of the clusters, and the late-type galaxies (LTGs) in the surrounding areas, would be a result partially due to the role played by environmental processes in the evolution of galaxies. In this sense, Wang et al. (2009) show that both early-type dwarf satellites and central galaxies could be affected by environmental effects in very similar ways, regardless of their locations within the cluster potential. This would indicate that the environment modifies the colour of the galaxy to a greater extent than its morphology.

The first catalogue of the Antlia cluster was provided by Ferguson & Sandage (1990, hereafter FS90) and consisted of 375 visually identified galaxies with a limiting magnitude of $B = 18$ mag covering an area of 6.1 square degrees. Due to the lack of spectroscopy, they also provided a membership status, mainly based on morphological criteria (Sandage, Binggeli & Tammann 1985), assigning status 1, 2, or 3 to ‘definite’, ‘likely’, or ‘possible’ members, respectively. The two brightest ellipticals (Es) in Antlia (NGC 3268 and NGC 3258), located in the central region, are considered as the dominant galaxies of the cluster. This cluster exhibits a high fraction of ETGs (considering those with membership status 1 and 2): $(Es + dEs + S0s + dS0s) = (0.05 + 0.59 + 0.07 + 0.05) = 0.76$, comparable with the Virgo and Fornax galaxy clusters, though Antlia has a central galaxy density a factor of 1.4–1.7 higher than them (FS90). By means of X-ray observations, it has been shown that Antlia is the nearest isothermal non-cool core cluster ($kT \sim 2$ keV, Wong et al. 2016), without a central brightness excess (Nakazawa et al. 2000). NGC 3268 is the centre of the X-ray emission, which is elongated towards a subgroup centred on NGC 3258 (Pedersen, Yoshii & Sommer-Larsen 1997; Nakazawa et al. 2000; Smith Castelli et al. 2008a). In addition, the galaxy-projected distribution presents a similar orientation, mainly in the direction defined by the two dominant Es. The Antlia cluster seems to be dynamically younger than Virgo and Fornax but unexpectedly relaxed when it is compared to other groups (Wong et al. 2016), which is also suggested by optical surveys (FS90; Hess et al. 2015). The globular cluster systems of the two brightest Es have been studied by Caso, Bassino & Gómez (2017, and references therein).

In the present paper, we perform a new surface photometry of the Antlia ETGs that complements previous papers (Calderón et al. 2015, hereafter Paper I, and Calderón et al. 2018, hereafter Paper II). Our study covers the largest area considered by any survey of this cluster until now, which motivates us to analyse the projected spatial density of galaxies up to ~ 700 kpc from the cluster centre, assuming an Antlia distance of ~ 35.2 Mpc (Dirsch, Richtler & Bassino 2003). We also recalculate the relations between structural parameters of the ETGs, now including more Es at the bright end than previous works. Moreover, we study the correlation between colour, effective radius, and radial velocity with the projected spatial distribution. Finally, we build the luminosity function (LF) of the cluster and estimate stellar masses.

This paper is organized as follows: in Section 2, we describe the data reduction, fit of the surface brightness profiles, membership classification, and the catalogue presented in this paper. In Section 3, we report the results on the new colour–magnitude relation (CMR), structural parameters, spatial projected distribution, and galaxy colour correlations. The discussion of the results is given in Section 4, and finally, the main conclusions are summarized in Section 5.

2 OBSERVATIONS AND DATA REDUCTION

We perform a photometric study on four new fields that complements those carried out in previous surveys on the Antlia cluster of galaxies. These fields were observed with the MOSAIC II camera mounted on the 4-m Blanco telescope at the Cerro Tololo Inter-American Observatory (CTIO, Chile), during the nights of 2010 February 11 and 12. They correspond to the ones labelled as 4–7 in Table 1, while the fields 0–3 are included in the same table for completeness, as they were studied in Paper I and obtained with the same instrumental setup. As in our previous studies, images were acquired selecting the Washington C and Kron–Cousins R filters; observations were later transformed into the genuine Washington C and T_1 bands (taking into account the zero-point offset: $R - T_1 = -0.02$). The coverage of MOSAIC II images is about 8000×8000 pixels with a scale of 0.27 arcsec pixel $^{-1}$ that corresponds to 36 arcmin \times 36 arcmin or, equivalently, to 370×370 kpc 2 , according to the adopted Antlia distance. Each field corresponds to a single MOSAIC II image, obtained by combining between five and seven dithered frames. Table 1 presents the basic information of each field including the date of observation, position of the field centre, filter, number of exposures (n_f) that were averaged to obtain the final image for each field, individual exposure time, mean airmass (X_R and X_C), and the seeing on the final R -band image. The image reduction was performed using the MSCRED package from IRAF, with the standard procedure already explained in the previous papers (e.g. Paper II). Regarding the calibration to the standard system, we obtained the following transformation equations for the new fields (4, 5, 6, and 7) to transform instrumental magnitudes (m_R and m_C) to standard ones (T_1 and C), based on standard stars fields from Geisler (1996):

$$\begin{aligned} T_1 &= (m_R + 0.02) + a_1 + a_2 X_R + a_3 (C - T_1) \\ C &= m_C + b_1 + b_2 X_C + b_3 (C - T_1), \end{aligned} \quad (1)$$

where the coefficients and their errors are given in Table 2. For the rest of the fields, we used the corresponding calibration equations given by Dirsch et al. (2003) for the central field, and Paper I for fields 1, 2, and 3. In addition, we applied a zero-point magnitude offset to refer the photometry to the central field, as explained in Paper I. The offsets in the T_1 band are between ~ 0.06 mag and ~ 0.1 mag, while in the C band, they are ~ 0.2 mag. All magnitudes and colours

Table 1. Observational data for the MOSAIC II fields (see text).

Field	Observation date	α_{2000}	δ_{2000}	Filter	n_f	Exposure (s)	Airmass	FWHM (arcsec)
0	April 2002	10:29:22	−35:27:54	<i>R</i>	5	600	1.059	1.0
				<i>C</i>	7	600	1.037	1.1
1	March 2004	10:28:59	−34:57:40	<i>R</i>	5	600	1.588	1.0
				<i>C</i>	7	1000	1.076	1.0
2	March 2004	10:31:09	−34:55:59	<i>R</i>	5	600	1.056	1.0
				<i>C</i>	7	900	1.016	1.2
3	March 2004	10:31:35	−35:30:42	<i>R</i>	5	600	1.269	0.9
				<i>C</i>	7	900	1.030	0.8
4	February 2010	10:27:01	−35:30:42	<i>R</i>	5	600	1.106	1.5
				<i>C</i>	6	900	1.025	1.9
5	February 2010	10:31:39	−36:06:59	<i>R</i>	5	600	1.278	1.1
				<i>C</i>	6	900	1.023	1.6
6	February 2010	10:29:38	−35:56:19	<i>R</i>	5	720	1.259	1.5
				<i>C</i>	7	900	1.080	1.6
7	February 2010	10:26:58	−35:55:22	<i>R</i>	5	600	1.025	1.2
				<i>C</i>	6	900	1.503	1.6

Table 2. Coefficients (top row) and errors (lower row) for the calibration equations 1.

a_1	a_2	a_3	b_1	b_2	b_3
0.466	−0.063	0.031	0.029	−0.405	0.096
0.026	0.018	0.007	0.059	0.052	0.007

depicted in the figures shown in this paper have been corrected for Galactic absorption and reddening. We used the colour excess $E(B - V)$ from Schlafly & Finkbeiner (2011), and the relations $E(C - T_1) = 1.97E(B - V)$ (Harris & Canterna 1977) and $A_R/A_V = 0.75$ (Rieke & Lebofsky 1985) in order to apply the corrections corresponding to the Washington photometric system.

2.1 Photometry of the Antlia cluster

The first paper presenting a CCD study of the Antlia cluster by Smith Castelli et al. (2008a) was based on the original Antlia catalogue (FS90), which involved 375 galaxies from early- to late-type morphologies. In Fig. 1, we compare the spatial distribution of the galaxies included in the FS90 catalogue (open circles) with the sample studied in the present work (open triangles). We labelled each MOSAIC II field from ‘field 0’ (the central one) to ‘field 7’. Their central coordinates are also given in Table 1.

The central field (0) was used in Smith Castelli et al. (2008a) to study 100 ETGs from the FS90 catalogue through isophotal photometry, while the existence of two compact elliptical (cE) galaxies in that sample was confirmed by Smith Castelli et al. (2008b). Then, new dE candidates (not previously catalogued) and radial velocities to confirm membership were added in Smith Castelli et al. (2012). In Papers I and II, we extended the galaxy sample by adding fields 1–3, calculating new total magnitudes and colours – against just isophotal photometry – of 107 additional cluster members plus 31 candidates with high membership probability. Their structural parameters, obtained by the fits of Sérsic models (Sérsic 1968), were also studied, and the corresponding scaling relations derived and compared with nearby clusters. In the present work, we add the surface photometry of galaxies in fields 4–7, so that our final catalogue contains 447 galaxies (332 of them considered as members

or candidates), which constitutes the largest sample of galaxies in the Antlia cluster with homogeneous integrated photometry.

The surface photometry analysis that we used has been extensively explained in Papers I and II. Here, we just include the main formulation applied to obtain magnitudes and structural parameters. We obtained the ETG surface brightness profiles through ELLIPSE task within IRAF and then fitting each of them with a one component Sérsic model:

$$\mu(r) = \mu_0 + 1.0857 \left(\frac{r}{r_0} \right)^{1/n}, \quad (2)$$

where μ_0 is the central surface brightness, r_0 is a scale parameter, and n is the Sérsic index, which is a measure of the concentration of the profile. From these parameters, we obtained the effective radius (r_e), effective surface brightness (μ_e), and total magnitudes and colours of the galaxy sample. Each profile was fitted using the task NFIT1D from IRAF, which implements the χ^2 statistic test by the Levenberg–Marquardt algorithm. In most cases, the residuals between model and observed profile turned out to be smaller than 0.5 mag (Paper II). Finally, as the parameters obtained from the fits were not affected by seeing for models with $n \lesssim 3$ (which is the general case), we did not apply further corrections (Paper II).

With the aim of making a complete catalogue of galaxies present in our images, we also included the surface photometry of all LTGs found on them. In these cases, we estimate a total magnitude, but the one-component scale parameters derived may differ from those of a more complex fit, with different components involved. In any case, LTGs were not included in the fits, regressions, or statistical tests (unless explicitly indicated).

2.2 Membership status of the ETGs

The membership status applied by FS90 was mainly based on morphological criteria due to the lack of spectroscopic data. Afterwards, it was proved with radial velocities that the FS90 status is in fact highly reliable (Smith Castelli et al. 2012; Calderón et al. 2015). In our series of papers on the Antlia cluster, the radial velocity range we have considered to define confirmed members is 1200–4200 km s^{−1} (Smith Castelli et al. 2008a). We have radial velocity measurements for 50 per cent of the ETGs with membership status 1 and 27 per cent status 2; out of them, the 95 and 82 per cent

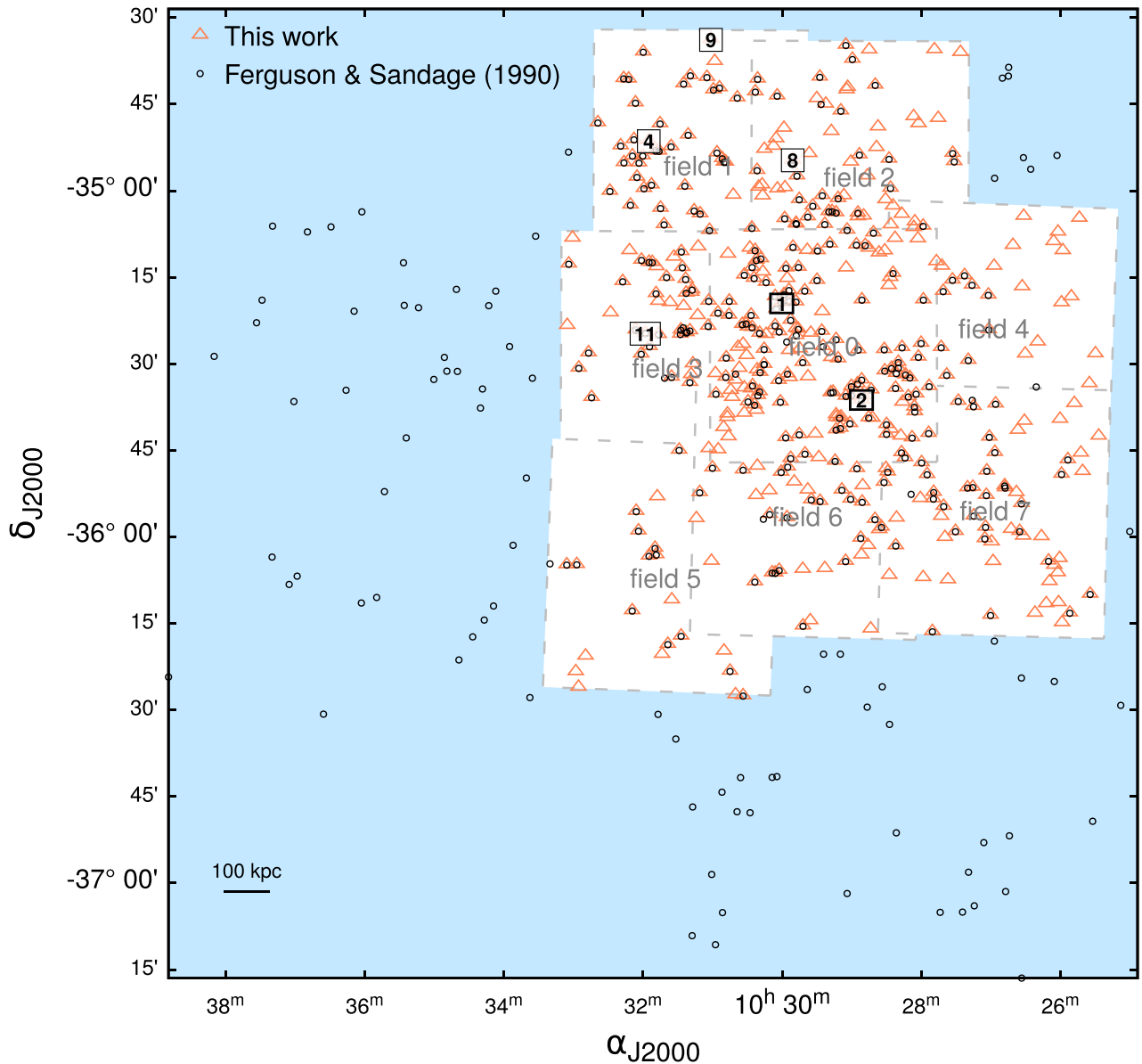


Figure 1. Position of the galaxies from the original [FS90](#) catalogue (open circles) and those studied in the present paper (from [FS90](#) and newly discovered ones, open triangles). Numbers 1 and 2 (within field 0) identify the two brightest elliptical galaxies NGC 3268 and NGC 3258, respectively, and numbers 4, 8, 9, and 11 (outside field 0) identify other galaxies brighter than $T_1 = 20$ (see Table 4 for their identifications and main properties).

are confirmed members, respectively. In Fig. 2, we show a histogram to compare the three different membership criteria that were applied to ETGs: (i) status defined by [FS90](#), (ii) measured radial velocities (Caso & Richtler 2015), and (iii) position in the colour–magnitude diagram, i.e. considering as members those galaxies located within $\pm 3\sigma$ of the mean CMR defined by ETGs (see Fig. 3 and Section 3.1), being this method extensively used to identify cluster members (e.g. Lieder et al. 2012). From now on, we will refer to this latter criterion as ‘3SmC’: $\pm 3\sigma$ mean colour criterion.

The original [FS90](#) catalogue has 112 ETGs with membership status 1, 79 with status 2, and 90 with status 3. Out of this total of 281 ETGs, 9 of status 1, 11 of status 2, and 28 of status 3 fall outside our MOSAIC II fields. Additionally, the identifications [FS90](#) 203/206 represent the same galaxy, as well as [FS90](#) 269/270. Several [FS90](#) galaxies could not be found on the MOSAIC II images: [FS90](#) 11, 42, 74, 151; and there are five galaxies with doubtful coordinates (big

shifts from [FS90](#)): [FS90](#) 178, 275, 204, 235, 287. When analysing just the ETGs of the [FS90](#) sample on the MOSAIC II images, we found that out of the 103 galaxies catalogued as ‘definite’ members, more than 80 per cent follow the 3SmC criterion, while almost ~ 50 per cent of them have radial velocities that confirm them as cluster members. Comparatively, galaxies in the two remaining membership statuses maintain a similar relation between those being confirmed by radial velocities and by the 3SmC criterion. Considering the 68 ‘likely’ members (status 2), more than 65 per cent are members following the 3SmC criterion, while 25 per cent have spectroscopic confirmation. Lastly, out of the 62 ‘possible’ members, 40 per cent follow the 3SmC criterion, and only about 5 per cent are spectroscopic confirmed members. These results (see also Fig. 2) reinforce the validity of the [FS90](#) membership status.

According to the previous results, we can consider ETGs with [FS90](#) membership status 1 plus those confirmed spectroscopically

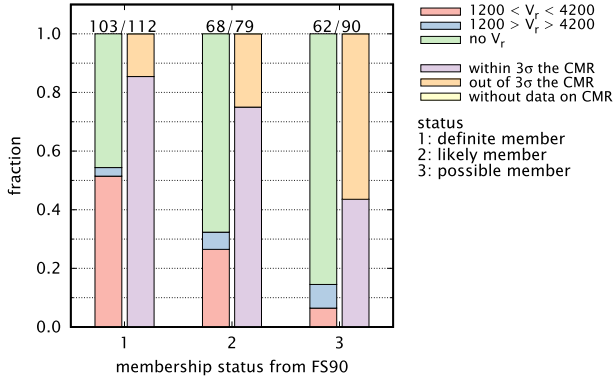


Figure 2. Comparison between the membership classification of ETGs: FS90 status, radial velocities, and ‘ $\pm 3\sigma$ from the CMR’ criterion. For each membership status, the right column corresponds to the CMR criterion and the left column to the radial velocity measurements.

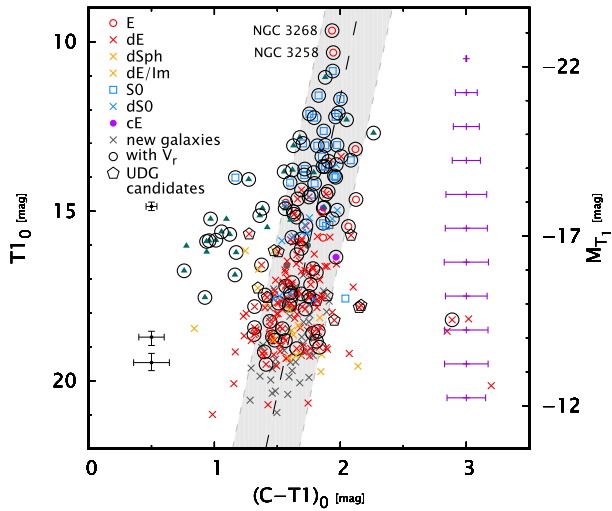


Figure 3. Colour-magnitude diagram for the ETGs in the final sample (i.e. confirmed members plus candidates following the 3SmC criterion). The grey band represents $\pm 3\sigma$ from the fitted CMR and colour codes for the symbols are given in the plot. On the left, we show errors in colour and magnitude. On the right, the colour dispersion of the ETG final sample for each magnitude bin. Additionally, we include LTGs identified with green triangles.

as fully confirmed members. In order to analyse quantitatively if the ETGs with membership status 2 can also be included in that sample of confirmed members (Buzzoni et al. 2012; Venhola et al. 2017), we calculated their residuals as the difference in magnitude and colour between their positions in the colour-magnitude diagram with respect to the mean CMR. Then, we performed a Kolmogorov–Smirnov (KS) test (Press et al. 1992) to compare the distribution of such residuals between two samples: the 17 ETGs with FS90 status 2 that are confirmed as members through radial velocities on one side, and the whole sample of 68 ETGs with FS90 status 2 on the other. As a result of the test, we obtained a KS statistic of 0.25 and a probability $1 - p = 0.68$. Thus, we cannot reject the hypothesis that the statistical distributions of the residuals of both samples, in the colour-magnitude diagram, come from a common distribution. On the basis of these results, we decided to add the ETGs with FS90 membership status 2 to the sample of ‘confirmed members’.

Additionally, we calculated the colour residuals with respect to the CMR for all ETGs in the confirmed member sample (e.g.

Hamraz et al. 2019) and fitted a Gaussian distribution to the resulting histogram (not shown). Then, we filtered out seven galaxies lying further than $3\sigma_{CT_1}$ of the Gaussian mean (with $\sigma_{CT_1} = 0.17$). Although these galaxies are confirmed members, their colour profiles seem to have unexpectedly large errors, so we decided not to take them into account for any further analysis.

2.3 The catalogue

The total number of galaxies identified in all fields is 447 (including FS90 and newly discovered galaxies of all morphologies). Out of them, we have obtained surface photometry for 300 (130 are new ones, uncatalogued in previous papers), which meet any of the following requirements: as ETG and LTG ‘confirmed members’, they have been confirmed by radial velocity or assigned FS90 membership status 1 or 2, or as just ETG ‘candidates’, they follow the 3SmC criterion (see Paper I). In Table 3, we present their photometric and structural parameters, with an asterisk indicating those galaxies that have been studied for the first time in the present paper.

Finally, there are 265 ETGs included in the catalogue, which we will call the *final sample*, i.e. ETG confirmed by radial velocity or assigned FS90 membership status 1 or 2, plus ETG ‘candidates’ that follow the 3SmC criterion.

3 RESULTS

3.1 Colour-magnitude relation

Fig. 3 shows the $(T_1)_0$ versus $(C - T_1)_0$ colour-magnitude diagram of the ETG final sample from our catalogue. We also included the LTGs represented by green triangles that have not been taken into account for any further statistics. The black dashed line shows the mean CMR calculated through a least-squares fit of the ETG ‘confirmed members’, taking into account errors in both axes, giving:

$$(T_1)_0 = (-17.5 \pm 0.2) (C - T_1)_0 + (46.7 \pm 3.3). \quad (3)$$

The standard deviation of the fit is $\sigma = 1.54$; the $\pm 3\sigma$ region about the CMR has been represented as a grey band in Fig. 3. Almost all ETGs confirmed as members by radial velocities, identified in Fig. 3 with black open circles, fall within $\pm 3\sigma$ from the mean CMR. The main characteristics of the CMR defined by ETGs in the Antlia cluster have been already discussed in Paper II. Now, with the addition of galaxies located in the new fields, we fill in the brighter side of the CMR and are able to see more clearly the ‘break’ at $M_{T_1} \sim -20$ mag (Jiménez et al. 2011), from which brighter galaxies keep an almost constant colour. Moreover, as most of the galaxy population lies within $\pm 3\sigma$ of the fit, the colour dispersion remains almost constant towards the faintest galaxies, as shown on the right side of Fig. 3. Twelve UDG candidates have been identified through their particular structural parameters (see next section for a detailed description).

3.2 Structural parameters

Using the data obtained from the new MOSAIC II fields (4–7), we have revisited the relations between structural parameters of the ETG population that involve central surface brightness, absolute magnitude, and Sérsic index. Following the same procedure as in Paper I and references therein, we obtained least-squares fit for the linear regressions of μ_0 versus M_V (Fig. 4a) and M_V versus n (Fig. 4b). The linear fits were obtained taking into account numerical

Table 3. Catalogue of galaxies in the Antlia cluster. The columns are described as follows: (1) id from [FS90](#) or new IAU identification, (2)–(3) J2000 coordinates, (4) morphology from [FS90](#), (5) membership status (from [FS90](#) except those with radial velocity), (6) Galactic extinction from Schlafly & Finkbeiner (2011), (7)–(12) global properties calculated in this work: T_1 -band magnitude ($C - T_1$) colour, T_1 -band effective surface brightness, effective radius, Sérsic index, and stellar mass (see Section 3.6), and (13) radial velocity (from Smith Castelli et al. 2012; Caso & Richtler 2015; Jones et al. 2009 or NED^a).

id	α_{J2000}	δ_{J2000}	Morph.	Mem	$E(B - V)$ (mag)	T_1 (mag)	$(C - T_1)$ (mag)	μ_e (mag arcsec ⁻²)	r_e (kpc)	n	$\log(M_G)$ (M_\odot)	V_R (km s ⁻¹)	Notes
1	10:25:05.040	−35:58:58.800	SmV	1	—	—	—	—	—	—	—	3220 ± 05	FS--
2	10:25:10.800	−36:29:09.600	S or ImV	3	—	—	—	—	—	—	—	—	F--
ANTLJ102532.4-354233.7	10:25:32.424	−35:42:33.721	S0	—	0.065	15.55	2.04	21.98	0.92	1.17	11.19	—	---C*
3	10:25:33.600	−36:49:15.600	dE	3	—	—	—	—	—	—	—	—	F--
ANTLJ102536.1-352816.3	10:25:36.115	−35:28:16.311	dE	—	0.066	18.83	2.45	22.61	0.27	1.14	9.95	—	---C*
4	10:25:38.147	−36:09:56.289	ImV or S	3	0.068	16.49	1.19	22.53	0.79	1.08	10.46	—	F--C*
ANTLJ102540.7-353326.6	10:25:40.718	−35:33:26.610	dE	—	0.062	18.57	2.36	22.72	0.37	0.60	10.05	—	---C*

Note. — The last column gives the references of the papers where the galaxy has been studied. The code FSCC refers to: (F) Ferguson & Sandage (1990); (S) Smith Castelli et al. (2008a) and Smith Castelli et al. (2012); (C) Caso & Richtler (2015); and (C) Calderón et al. (2015) and Calderón et al. (2018). Asterisks identify galaxies studied for the first time in the present paper. The full table can be accessed electronically.

^aThe NASA/IPAC Extragalactic Database (NED) is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

errors in both axes and excluding the cEs of the sample. The resulting relations are:

$$\mu_0 = (39.41 \pm 1.03) + (1.19 \pm 0.06) M_V \quad (4)$$

$$M_V = (-13.75 \pm 0.88) + (-11.29 \pm 4.21) \log(n). \quad (5)$$

By combining equations (4) and (5), we obtain the following linear relation between μ_0 and n (Fig. 4c):

$$\mu_0 = 23.04 - 13.43 \log(n). \quad (6)$$

Based on these linear fits, we have rebuilt the relations that involve the effective parameters: effective surface brightness (μ_e) and effective radius (r_e), which are shown in Figs 4(d) and (e), respectively.

In [Paper I](#), we compared the structural parameters of the Antlia galaxy population with those of other clusters and groups. The recent releases of the Next Generation Fornax Survey (NGFS, Eigenthaler et al. 2018) and Fornax Deep Survey (FDS, Venhola et al. 2017) allow us to compare our result with those from ETGs of nearby clusters. The optical side of the NGFS is based on deep multiband observations (u' , g' , and i' bands), taken with the Dark Energy Camera (DECam, Flaugher et al. 2015), mounted on the 4-m Blanco telescope at the CTIO. The FDS survey, in turn, used the ESO VLT Survey Telescope (VST), a 2.6-m wide-field optical telescope at Cerro Paranal, Chile. The observations were taken in u , g , r , and i bands with the OmegaCam camera (Kuijken et al. 2002). Despite that NGFS achieved a larger coverage than FDS, both surveys span almost the same magnitude range: $-16 < M_V < -8$ mag (magnitude transformations were performed according to Fukugita, Shimasaku & Ichikawa 1995) and list in their respective catalogues the total magnitude, Sérsic index, and effective radius of each galaxy. From these quantities, we calculated μ_e using the following expressions:

$$\mu_0 = m + 1.995450 + 5 \log(r_e) \quad (7)$$

$$\begin{aligned} & -5n \log(b_n) + 2.5 \log(n \Gamma(2n)) \\ \mu_e &= \mu_0 + 1.0857 b_n, \end{aligned} \quad (8)$$

where $\Gamma(x)$ is the complete gamma function, and b_n can be obtained by solving $\Gamma(2n) = 2 \gamma(2n, b_n)$ (Ciotti 1991), where $\gamma(a, x)$ is the incomplete gamma function. In Figs 4(d) and (e), we have included data from the NGFS and FDS catalogues, with different symbols identifying the UDGs (Venhola et al. 2017), that have larger radii

than dwarf galaxies. The data from the Fornax samples extend the fainter limit of our Antlia data in about 3 magnitudes.

We have searched for UDGs in Antlia, taking into account their characteristics as observed in Virgo, Fornax, and Coma clusters (van Dokkum et al. 2015; Venhola et al. 2017). If we consider as UDG candidates all dwarf galaxies with $r_e > 1.5$ kpc, then the UDGs in Antlia attain a total number of 12 candidates. These galaxies have a mean effective radius (r_e) = 1.88 ± 0.35 kpc, and absolute magnitudes in the range $-13.9 > M_V > -16.4$ mag, being the mean value $\langle M_V \rangle = -15.0 \pm 0.9$ mag. They also have surface brightnesses $\mu_e > 23$ mag arcsec⁻², and a mean Sérsic index of $\langle n \rangle = 1.44 \pm 0.53$. This means that the UDGs candidates have larger radii and fainter effective surface brightness than the ETG population, at the same absolute magnitude, as can be seen in Figs 4(e) and (d), respectively. We will come back to this topic in the Discussion.

3.3 Spatial distribution

In order to study the projected spatial distribution, a position within the cluster should be taken as a reference point. Although the Antlia cluster has two dominant cD galaxies, to select such central point, we will take into account the intensity of the X-ray emission, according to the map given by Wong et al. (2016, see their fig. 1). It can be seen that the X-ray intensity is clearly higher at the position of the galaxy NGC 3268, while only an extension of lower intensity towards NGC 3258 is present. Thus, in the following, NGC 3268 will be considered as the centre of the projected spatial distribution. In addition, NGC 3268 is the centre of the youngest subcluster in Antlia (Hess et al. 2015). In Fig. 5, we present the position of the galaxies in our sample, superimposed to the X-ray map from Wong et al. (2016) that covers the central part of the cluster. The smoothed colour distribution shows the X-ray intensity (high: yellow–red, low: dark blue). The grey dashed line represents the boundaries of MOSAIC field 0, the symbols are as in the previous figures, and the white square and pie regions correspond to different X-ray observations. We have also indicated the position of galaxies brighter than $M_{T_1} = -20$ mag in Table 4. In this case, the locations of NGC 3268 and NGC 3258 are indicated with the upper and lower black crosses, respectively. Four of the brightest galaxies in our sample ([FS90](#) 172, 253, 300, and 304; or id 8, 9, 4, and 11, respectively) are located outside field 0, on field 1 (towards the north-east of field 0), showing that the brightest galaxies lie in the direction that connects both cD galaxies, coincident with the main axis of the elongated X-ray distribution (see Fig. 6).

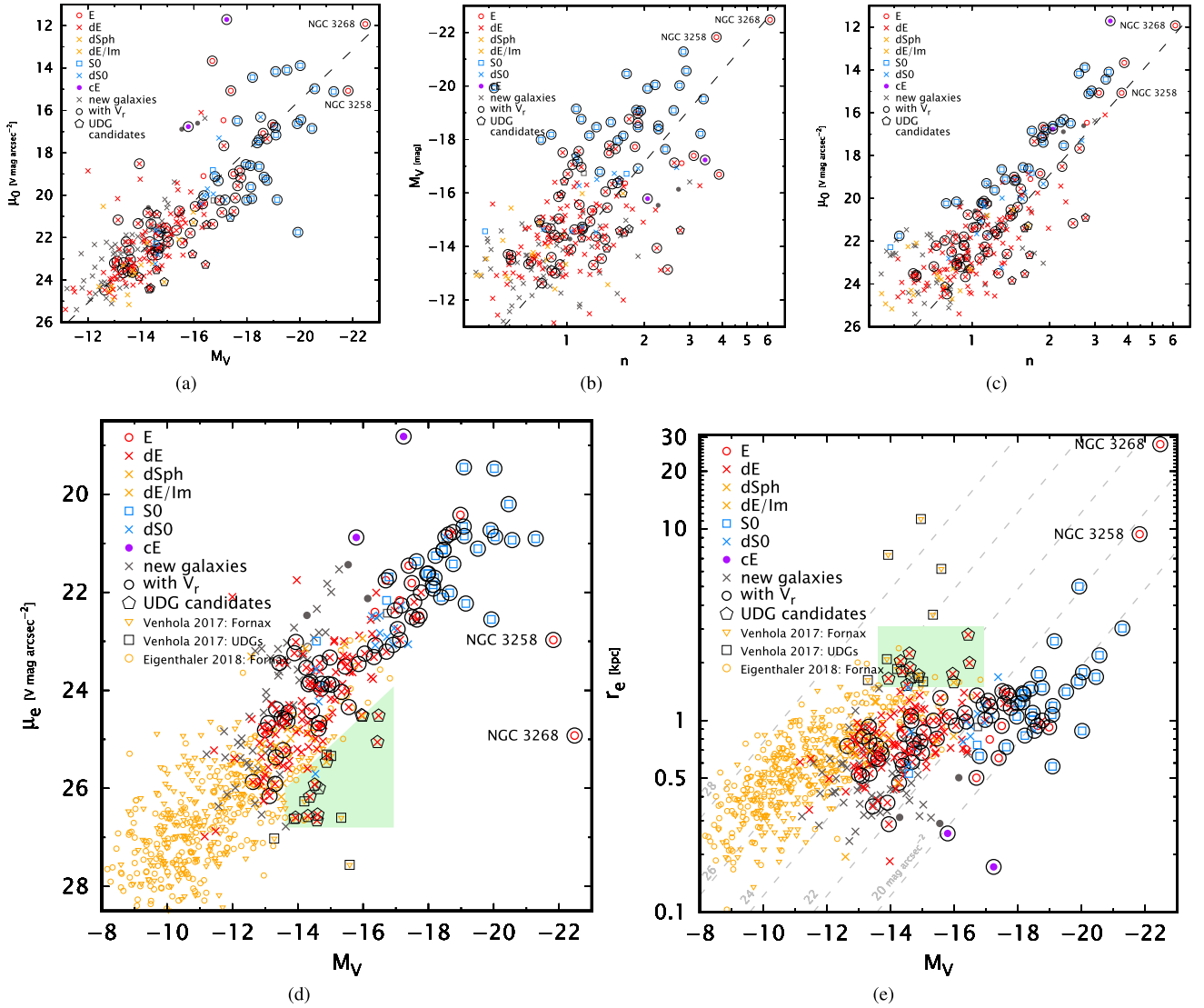


Figure 4. Relations between the structural parameters for the ETG final sample. Symbols are identified in each plot. Upper panels: (a) Central surface brightness versus absolute magnitude (V band). (b) Absolute magnitude versus logarithm of the Sérsic index (V band). (c) Central surface brightness versus logarithm of the Sérsic index (V band). Dashed lines show the respective least-squares linear fit, except in (c) that is obtained from (a) and (b) fits (see text). Lower panels: (d) Effective surface brightness versus absolute magnitude (V band). (e) Logarithm of the effective radius versus absolute magnitude (V band). With grey dashed lines, we indicate the loci of constant mean effective surface brightness ($\langle\mu_e\rangle$). Green areas in (d) and (e) highlight the position of Antlia UDG candidates.

The Antlia cluster-projected spatial distribution of the complete sample studied in this paper is shown in Fig. 6, including the eight MOSAIC II fields (numbered from 0 to 7). Taking into account the overlap between fields, the total area covered in this paper is about $100 \text{ arcmin} \times 100 \text{ arcmin}$ ($\sim 1000 \times 1000 \text{ kpc}^2$). Galaxy symbols are explained in the plot, numbers identify galaxies brighter than $T_1 = 20$ located outside field 0, and the dominant galaxies are also highlighted. We clearly indicate galaxies with radial velocity measurements (encircled in black), new uncatalogued galaxies (grey crosses), and UDG candidates (black pentagons).

In order to study any possible correlation between galaxy colour and projected spatial distribution, as well as for the following analysis, we split the ETG sample into galaxies bluer and redder than the CMR (equation 3). In Fig. 7(a), the global density profile shows that the main concentration of ETGs is clearly around NGC 3268 (i.e. cluster-centric radius $R = 0 \text{ arcmin}$), while another overdensity is present at the fourth annulus, close to NGC 3258. The galaxies

bluer than the CMR are more concentrated up to $R \sim 300 \text{ kpc}$, with a smooth decline. Particularly, in the innermost annulus, bluer ETGs more than double the redder ones. The overdensity close to the location of NGC 3258 is also detected in the galaxies redder than the CMR, and it corresponds to a younger subcluster that is probably infalling to NGC 3268 (Smith Castelli et al. 2008a; Hess et al. 2015).

We also studied the projected spatial distribution considering giant (bright) and dwarf ETGs. The former, mostly Es and S0s, were represented in Fig. 7(a) by empty squares, while the latter, mostly dEs, by empty circles. The adopted morphology for the FS90 galaxies was taken from their catalogue. The new galaxy candidates added to the final sample, mainly dwarfs, were selected by the elliptical shape on the images (see Paper I). Frequently, the distinction between dwarfs and giants is not clear. Rather than the magnitude, the central surface brightness is considered as a better option to separate dwarfs from giants (e.g. Barazza et al. 2009; Kormendy et al. 2009). We thus split the sample into dwarfs and

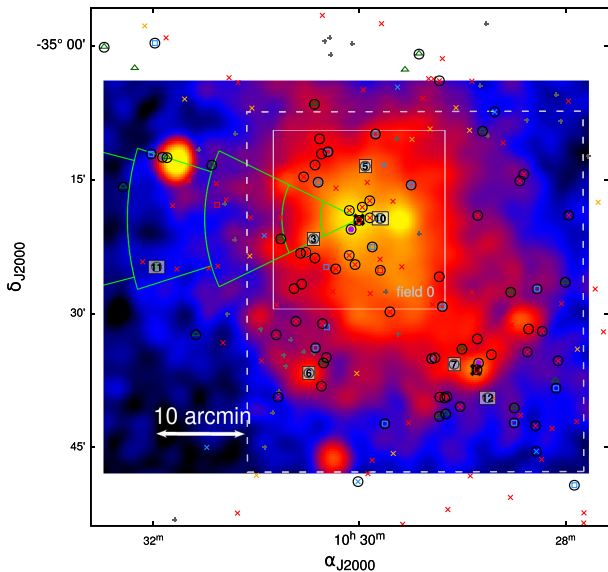


Figure 5. Central region of the Antlia cluster (field 0), with the location of the sample galaxies (symbols as in previous figures), superimposed on a smoothed X-ray ROSAT PSPC (0.5–2 keV band) map reproduced from Fig. 1 (upper left) of Wong et al. (2016) by permission of the AAS. X-ray intensity varies from high in yellow–red to low in blue–black. The grey dashed line represents the MOSAIC field 0, and the white square and pie regions correspond to different X-ray observations. Numbers identify bright galaxies listed in Table 4.

giants, taking as limit between them: a central surface brightness $\mu_0 = 19 \text{ mag/arcsec}^2$ (V-band), a magnitude $T_1 = 14 \text{ mag}$ (Smith Castelli et al. 2008a), or just the morphological classification. We then applied a KS test to compare their projected distributions (both versus cluster-centric radius or azimuthal angle), and all resulted comparatively quite similar, i.e. the KS test shows that we cannot distinguish statistically the different samples, when they are split by μ_0 , T_1 or the morphological classification. For the rest of the analysis, we will apply a dwarf/bright ETGs separation using μ_0 .

The projected density of the faint ETGs surpasses that of the bright ones in every annulus of the profile (Fig. 7a), even in the central region of the cluster, where the largest number of bright galaxies is found. To determine the principal direction along which the galaxies are distributed, we show the azimuthal-projected distribution in Fig. 7(b), which is also centred on NGC 3268. It can be seen that the main overdensities take place in the direction of NGC 3258 ($\theta = 0^\circ$) and in its exactly opposite direction ($\theta = 180^\circ$), i.e. an axis connecting both cD galaxies defines the direction along which galaxy density is enhanced. For the sake of clarity, the circular sectors corresponding to these overdensities are also shown in Fig. 6. The correlation between the X and Y coordinates of the galaxies can be quantified with the Pearson coefficient (Freedman & Purves 2007), which is about 0.2, a moderate correlation with a confidence level greater than 95 per cent. Also, the orientation of the correlation ellipse matches the direction of an axis joining NGC 3258 and NGC 3268. This indicates that we cannot reject the hypothesis that the projected distribution is elongated in the direction corresponding to the overdensities shown in Fig. 7(b). If we compare the projected distributions of dwarf against bright galaxies using a KS test, we obtain statistics of 0.277 and 0.118 ($1 - p$ values of 0.9868 and 0.2459) for radial and azimuthal distributions, respectively. Such results indicate that we can reject the hypothesis that the two populations are drawn from the same distributions, with a confidence level of 95 per cent, when we consider

the radial distribution, but we cannot reject such hypothesis for the azimuthal distribution centred on NGC 3268.

On the other hand, the radial and azimuthal distributions for the full final sample were also compared using the KS test, in order to look for any possible colour-projected spatial distribution trend. In this case, the KS test gives statistics of 0.1072 and 0.1004 (with $1 - p$ value of 0.5797 and 0.5144) for the radial and azimuthal distributions, respectively. This means that we cannot reject the hypothesis that the ETGs redder and bluer than the CMR come from the same parent distribution.

In Fig. 8, we investigate the relation of the ETGs ($C - T_1$)₀ colour (top panel) and effective radius (bottom panel) versus the cluster-centric radius (R , centred on NGC 3268). The galaxy sample is again divided into those bluer (blue triangles) and redder (red open circles) than the CMR regression, while filled circles show the respective mean values in each bin. Black open circles correspond to the mean values of all ETGs in each bin and the grey band indicates their dispersion. As a reference, the brighter galaxies have been identified in the respective bins, following the id given in Table 4. The colour of the whole final sample, as well as the blue and red subsamples (top panel), does not show any clear gradient along the cluster-centric distance but seems to be almost constant at $\langle (C - T_1)_0 \rangle \sim 1.69 \pm 0.09 \text{ mag}$ for the whole radial range (the correlation coefficient for the linear fit is $cor = 0.06$). One of the most populated bins corresponds to $R = 22.5 \text{ arcmin}$, which agrees with the position of NGC 3258 and has the largest dispersion $\sigma_{22.5} = 0.24 \text{ mag}$. This behaviour is the same if we refer to the bluer or redder subsamples, separately.

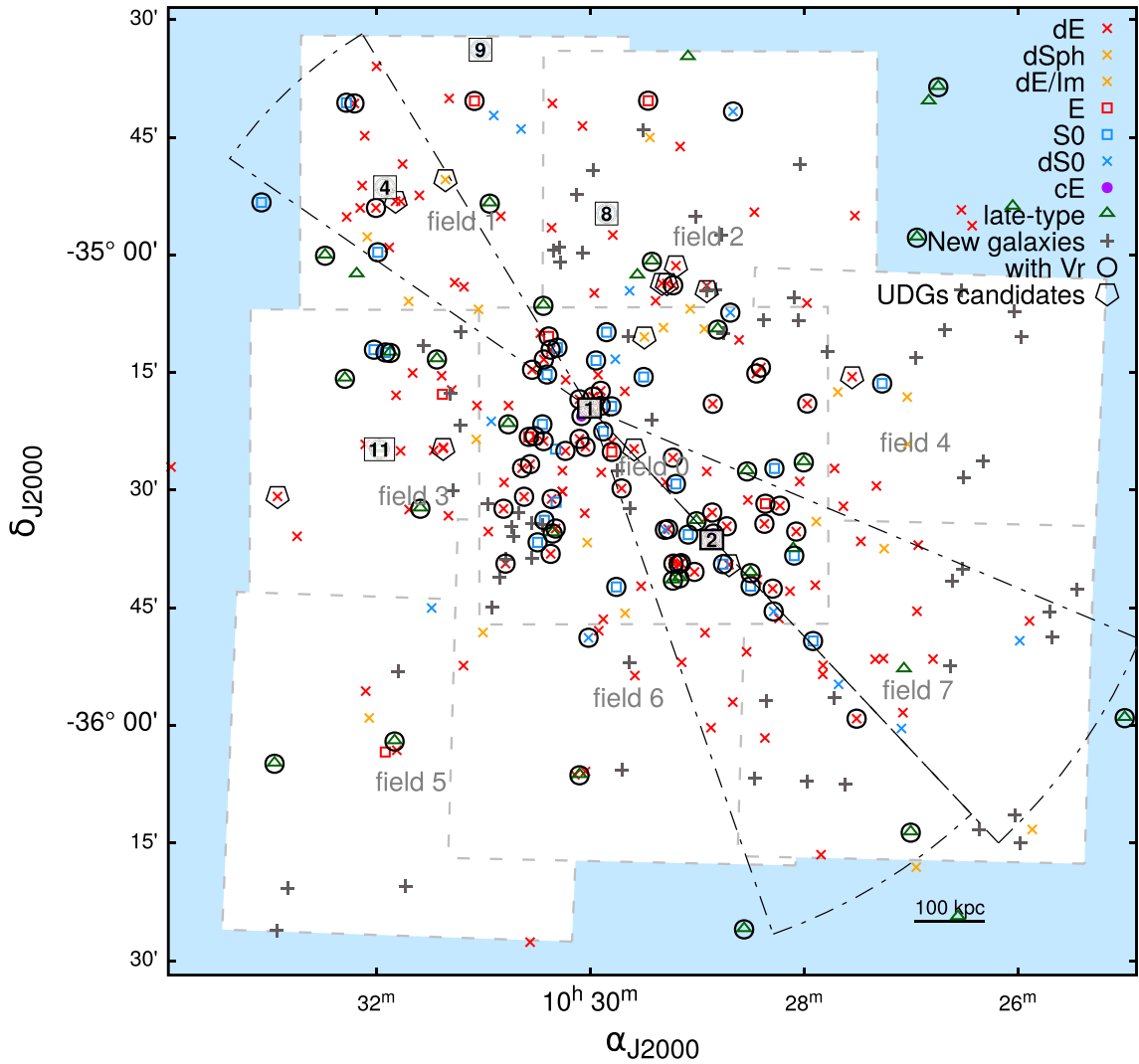
Now, we consider the radial trend of effective radius (bottom panel) of the whole ETG sample, excepting the following outliers: three bright galaxies with $r_e > 5 \text{ kpc}$ (NGC 3258, NGC 3268, and FS90 253) and the two cEs (FS90 110 and FS90 192). The mean effective radius close to the cluster centre (NGC 3268) is $\langle r_e \rangle = 0.9 \pm 0.2 \text{ kpc}$, which is mainly traced by the dEs (e.g. Smith Castelli et al. 2008a). A correlation ($cor = 0.9$ for a linear fit of r_e versus R) is present along the cluster-centric distance, in the sense that the mean effective radius decreases towards the outer regions, even beyond the subgroup centred on NGC 3258. The galaxies that present redder and bluer colours than the CMR follow the same trend along the cluster radius. In particular, at $R = 22.5 \text{ arcmin}$, the redder galaxies have the largest dispersion in mean effective radius $\sigma_{22.5} = 0.57 \text{ kpc}$, while if we consider the whole sample or the bluer galaxies, the largest dispersion is present at the next bin, at $R = 27.5 \text{ arcmin}$ with $\sigma_{27.5} = 0.52$ and 0.61 kpc , respectively.

3.4 Radial velocity distribution

In this section, we revisit the radial velocity distribution as a function of R . The radial velocities of galaxies in the Antlia Cluster core were extensively studied by Caso & Richtler (2015). We can now complement them with our photometry and structural parameters. In Fig. 9(a), we show the radial velocity histogram of the total number of galaxies with radial velocities lying within the membership range (grey) and split the sample into galaxies redder and bluer than the CMR. The three arrows indicate the locations of the main peaks in the histogram obtained by Caso & Richtler (2015): $2060 \pm 200 \text{ km s}^{-1}$, $2780 \pm 100 \text{ km s}^{-1}$, and $3600 \pm 130 \text{ km s}^{-1}$. It can be seen that the extreme radial velocity peaks seem to be related with the bluer ETGs, while red ETGs tend to be concentrated to the central peak (which includes the brightest galaxies) and whose location could be related with the X-ray major emission. On the other hand, the bluer galaxies tend to display extreme radial velocities

Table 4. Washington photometry of Antlia galaxies brighter than $M_{T_1} = -20$ mag. Radial velocities (last column) were taken from Caso & Richtler (2015).

id	FS90	NGC/IC	α_{J2000}	δ_{J2000}	$(C - T_1)_0$ (mag)	T_{10} (mag)	V_R (km s $^{-1}$)
1	185	3268	10:30:00.65	-35:19:31.60	1.93	9.67	2800 \pm 21
2	111	3258	10:28:53.29	-35:36:20.10	1.94	10.31	2792 \pm 50
3	224	3271	10:30:26.42	-35:21:34.40	1.93	10.86	3737 \pm 27
4	300	3281	10:31:52.00	-34:51:12.37	1.87	11.04	3200 \pm 22
5	184	3269	10:29:57.03	-35:13:27.57	1.82	11.58	3754 \pm 33
6	226	3273	10:30:29.09	-35:36:36.60	2.00	11.68	2660 \pm 08
7	125	3260	10:29:06.33	-35:35:42.45	1.98	12.10	2439 \pm 46
8	172	2584	10:29:51.46	-34:54:42.49	1.75	12.12	2549 \pm 19
9	253	2587	10:30:59.84	-34:33:44.52	1.96	12.21	2111 \pm 18
10	168	3267	10:29:48.58	-35:19:19.03	1.79	12.23	3709 \pm 33
11	304	3258D	10:31:55.70	-35:24:35.15	2.04	12.30	2476 \pm 16
12	105	3257	10:28:47.08	-35:39:28.69	1.86	12.62	3237 \pm 15

**Figure 6.** Composite image of the eight MOSAIC II fields with Antlia galaxies from the complete sample superimposed. Symbols are explained on the upper right corner and numbers 4, 8, 9, and 11 identify galaxies brighter than $T_1 = 20$ (listed in Table 4), located outside field 0. Both dominant galaxies are also labelled. The three circular sectors indicate high density of galaxies (see Fig. 7b and Section 3.3).

within the membership range and have a flat-projected distribution around the central field. Regarding the radial velocity distribution in this central region, we note that there are two S0 galaxies, NGC 3273 and NGC 3260 (FS90 226 and FS90 125, respectively),

located in areas of high X-ray emission between both cDs, that have the lowest radial velocities within the central field. On the other side, the highest radial velocities in this field correspond to NGC 3269 and NGC 3271 (FS90 184 and FS90 224) that lie

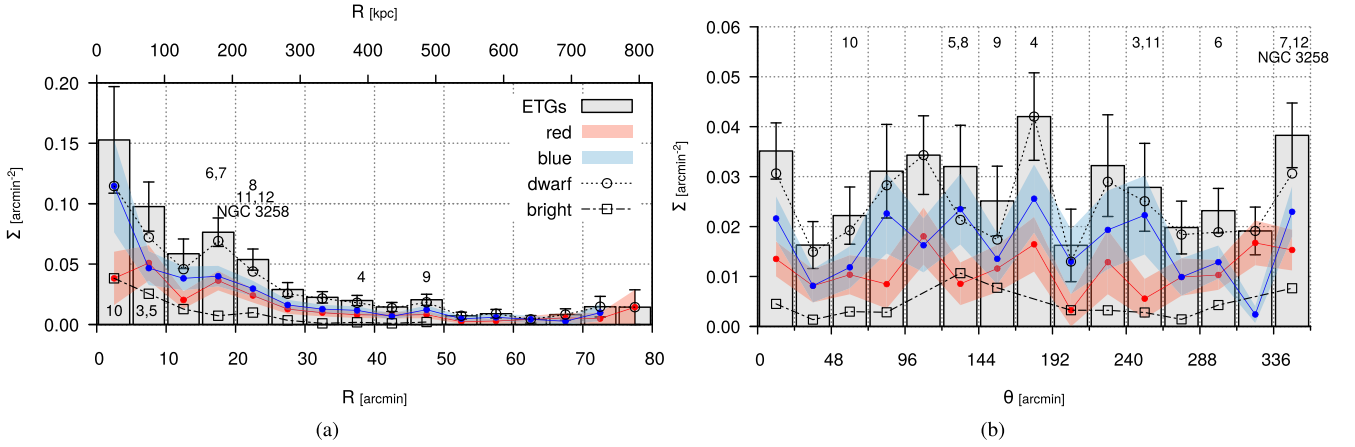


Figure 7. Projected number density of Antlia galaxies with respect to the centre of the distribution, taken at NGC 3268, for all ETGs in the final sample. The galaxies brighter than $T_1 = 20$ mag are indicated in the corresponding bin. Blue and red lines correspond to galaxies bluer and redder than the CMR, while open squares and circles indicate bright and faint ETG subsamples, respectively. Panel (a): Density profile of the projected number density Σ versus cluster-centric radius R . Panel (b): Azimuthal distribution of the projected number density, taking the direction to NGC 3258 as $\theta = 0^\circ$.

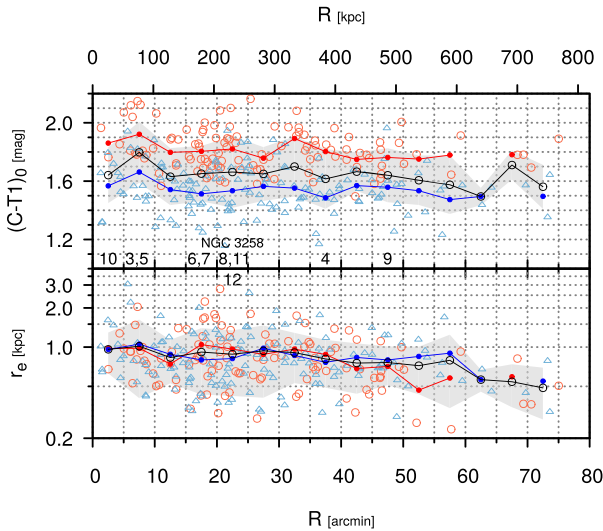


Figure 8. ETGs colour and effective radius as a function of cluster-centric distance, centred on NGC 3268. Galaxies bluer and redder than the CMR are identified with blue triangles and red circles, respectively. Filled blue and red circles correspond to the mean values in each bin. The mean values for all galaxies in each bin are shown with open black circles while their dispersion is represented by the grey band. Numbers between panels identify the galaxies brighter than $T_1 = 20$ mag (Table 4) in the corresponding bin, as in previous figures.

near NGC 3268 in the area of maximum X-ray emission of the cluster.

In Fig. 9(b), we represent the radial velocity distribution together with the spatial projected distribution of galaxies as a ‘heatmap’, i.e. a two-dimensional graphical representation that uses a colour palette to indicate the different velocity ranges, each one calculated as the distance weighted mean. We have decided to re-centre the V_R range relative to the NGC 3268 velocity, which is $\sim 2800 \text{ km s}^{-1}$. In this way, the yellow/green areas have a predominance of galaxies with similar V_R , red/orange and blue areas indicate extreme V_R (within the membership range). We also indicate the position of the brightest galaxies as spatial reference of the substructures in the cluster (ids in Table 4). The three kinematic substructures studied by

Hess et al. (2015): I, II, and III, are also indicated. These structures were obtained taking into account the velocity dispersion referred to a local mean velocity. Additionally, we performed a spatial clustering study through the k-means numerical method (Lloyd 1982). This algorithm is based on an Euclidean distance and iteratively calculates the distance between groups of objects, built from the average distance of each group. As a result, we can determine the centroid of each data group in the space of coordinates (in our case, the spatial projected coordinates). We show in Fig. 9(b), with filled triangles (pointed at by arrows), the position of the centroids when two groups are considered as substructures, taking into account the members of the sample with V_R . Finally, in Figs 9(c) and (d), we show separately the positions of the galaxies redder and bluer than the CMR.

3.5 Luminosity function

The LF of a cluster can be built by counting the number of galaxies in different magnitude bins. We show in Fig. 10 the LF, in the T_1 band, of the ETGs in the Antlia Cluster. We used the 265 ETGs in the sample, which are distributed on the eight MOSAIC II fields, covering an effective area of about 2.6 deg^2 or 1 Mpc^2 at the distance of 35.2 Mpc ($H_0 = 75 \text{ km s}^{-1} \text{ Mpc}^{-1}$, Dirsch et al. 2003). We used the Schechter function (Schechter 1976),

$$\varphi(M)dM = \varphi^* 10^{-0.4(M-M^*)(\alpha+1)} \exp^{-10^{-0.4(M-M^*)}} dM, \quad (9)$$

where φ^* is the characteristic density, M^* is the absolute magnitude at the slope break, and α is the slope at the faint luminosity end. We show in dashed-black line a fit of the LF in the magnitude range $-22 < M_{T_1} < -14$ mag, with parameters: $\varphi^* = 0.55 \pm 0.36$, $\alpha = -1.37 \pm 0.03$, and $M^* = -21.70 \pm 1.36$, and $\chi^2 = 0.11$.

Due to the scarcity of bright (giant) ellipticals in the cluster, the Schechter fit cannot be well constrained at the bright magnitude side. The faint end is affected by the sample incompleteness. Since the sample was visually selected, automatic methods to calculate the completeness cannot be applied (Lieder et al. 2012). To estimate the incompleteness, we used the structural characteristics of the faintest galaxies in the sample, determining the lowest effective surface brightness detected for a typical galaxy, which turned out to be about 2 per cent at $M_V \sim -13$ mag (Paper II). The absence

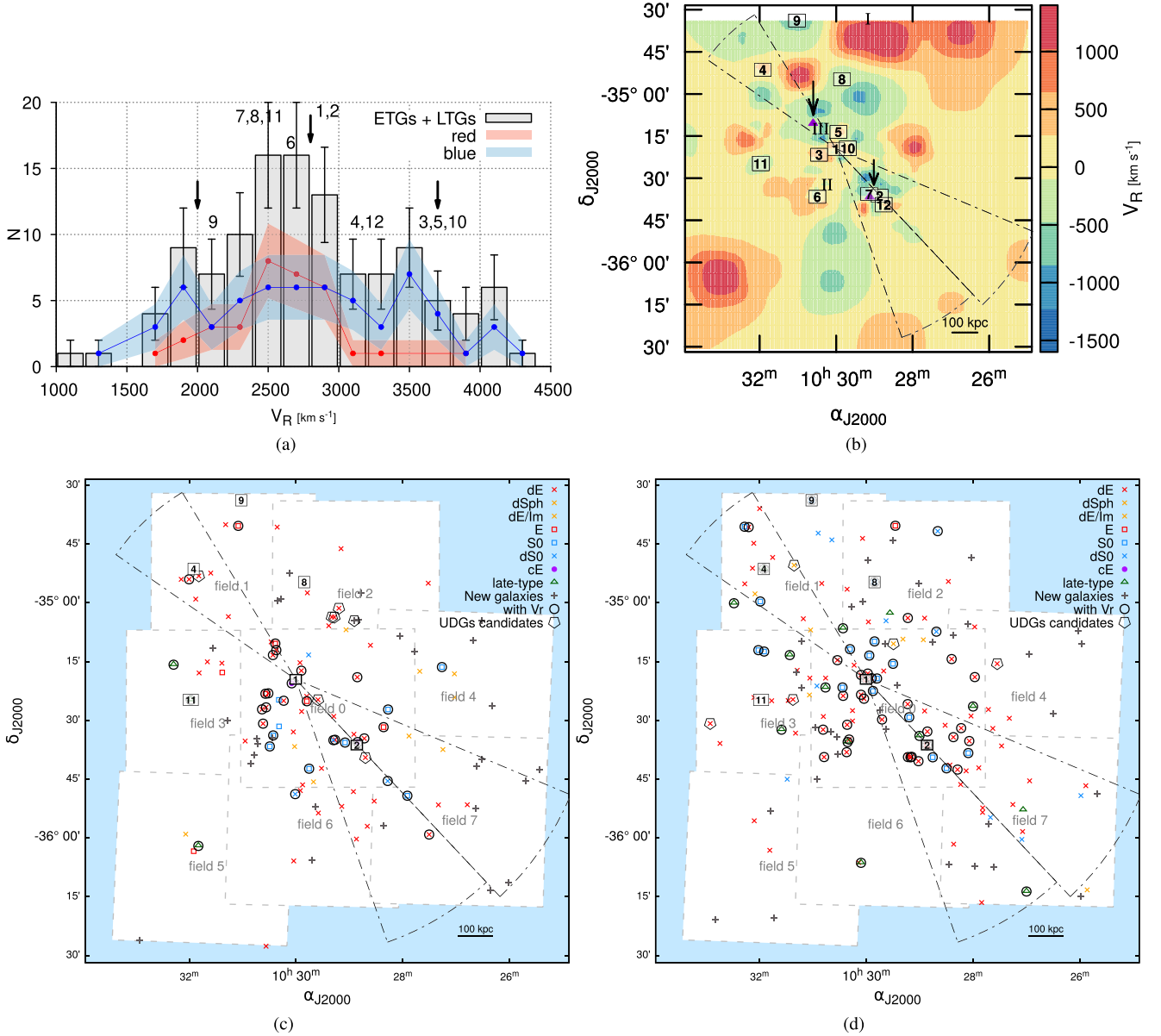


Figure 9. (a) Radial velocity histogram of all galaxies in the sample split into redder and bluer than the CMR. With black arrows, we indicate the main peaks in the histogram obtained by Caso & Richtler (2015) and identify the brightest galaxies in the respective bins, following the id given in Table 4. (b) The Antlia cluster radial velocity map and the relation with the projected spatial distribution. With filled triangles (highlighted with arrows), we have shown the position of the two centroids (see text). (c) and (d) The projected spatial distributions of redder and bluer galaxies than the CMR are shown, respectively. We also show on the maps the three circular sectors with the highest galaxy densities.

of a dip in the LF of Fig. 10, which was extensively studied in different clusters and environments (and references therein Lee et al. 2016), agrees with similar results in Fornax (Hilker, Mieske & Infante 2003; Jordán et al. 2007) and the NGC 5044 group (Buzzoni et al. 2012).

3.6 Estimation of stellar masses

In order to estimate the galaxy stellar mass of each object in the sample, we used the galaxy models presented in Buzzoni (2005) that provide the appropriate mass-luminosity ratio for different galaxy morphologies. The galaxy luminosity in the T_1 band was transformed into V using the expression given in Buzzoni (2005, table A1 and A2) and the M/L ratios listed in Buzzoni et al. (2012, Table 4).

The galaxy mass (M_G) was estimated as follows:

$$M_G = \frac{M}{L} \left[\frac{M_\odot}{L_\odot} \right] 10^{-0.4(V-V_\odot)}. \quad (10)$$

In Fig. 11, we show a stellar mass histogram for the 265 ETGs in the sample. All the galaxies span the range from $7.5 \lesssim \log(M_G)/M_\odot \lesssim 14.0$, with an abrupt cut at the lower mass limit. The logarithmic mean mass is $10.48 \pm 0.78 M_\odot$, which corresponds to $\langle M_G \rangle = 3.0 \times 10^{10} M_\odot$. We also show the distribution of blue/red galaxies (with respect to the CMR) in the same figure. Blue galaxies tend to be more frequent along almost all the mass range, while for intermediate masses ($\log(M_G)/M_\odot \sim 10.0 - 11.5$), redder galaxies show a deficit with respect to the general trend.

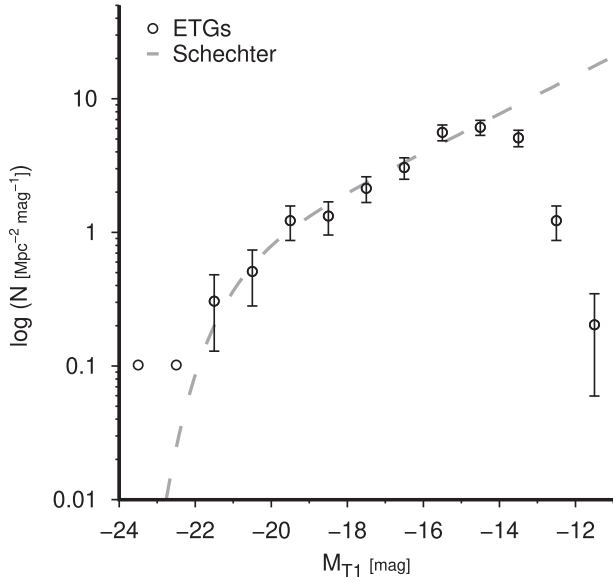


Figure 10. Luminosity function of the ETGs in the sample and a Schechter function fit on dashed line.

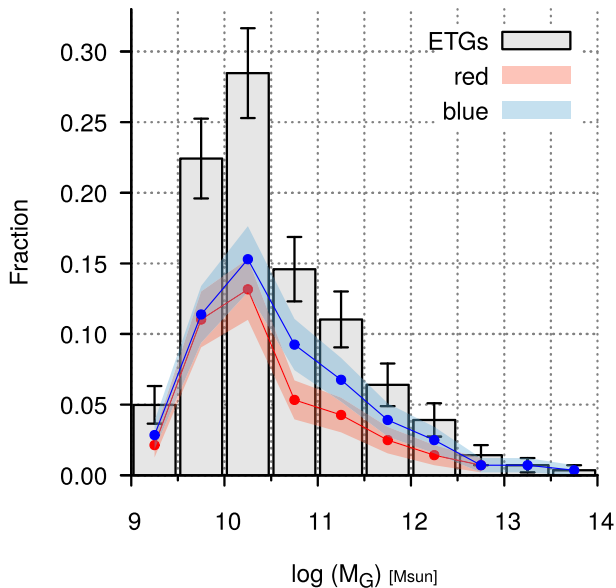


Figure 11. Galaxy stellar mass histogram.

4 DISCUSSION

In this work, we study the largest ETG sample of the Antlia Cluster. Our main goals are to quantify the characteristics of the ETG population and to describe their projected spatial distribution as well as the correlations between their structural parameters. Table 3 consists of 543 galaxies from all types, including those identified by FS90, Smith Castelli et al. (2012), and Paper I. Each source is clearly identified in the table.

4.1 Structural parameters

Fig. 4 presented the relations between the structural parameters of all 265 ETGs in the sample. Each parameter: r_e , Sérsic index n and μ_e , was obtained from the fit of the respective surface brightness profile;

their relationships reveal the different populations that coexist in the cluster. Now, in Fig. 12, we show the distributions of each parameter along with that of the total absolute magnitude M_V (for the maths, see Paper II). On each graph, we show in grey the histogram of the ETGs in the sample, in red and blue the corresponding subpopulations with respect to the CMR, and we additionally show them separated into giant (empty squares) and dwarf (empty circles) ETGs, using as limit the central surface brightness $\mu_0 = 19$ mag arcsec $^{-2}$ (on V band). Fig. 12(a) shows the distribution of the total absolute magnitude, spanning a range of $-23 < M_V < -10$ mag. The LSB galaxies dominate the faint magnitude range, with a maximum at $M_V \approx -14$ mag. This behaviour is followed by both the red/blue galaxies, with the exception that the bluer galaxies seem to dominate in number at the faintest magnitudes. Fig. 12(b) shows a unique peak at $r_e \lesssim 1$ kpc, shared by the blue/red and dwarf/giant subpopulations. While small radii bins are populated mainly by dwarf galaxies, giant ones span a broad range in radius. We remind here that we use a one-component fit to model the surface brightness profile, which is very accurate for most of the sample galaxies. A sudden drop can be seen at $r_e \approx 1.5$ kpc, which sets the radius where LSB galaxies start to be dominated by UDGs, as those found in Fornax (Venhola et al. 2017; Eigenthaler et al. 2018). If we take into account the dwarf population, the characteristic effective radius is $r_e = 0.8 \pm 0.4$ kpc (T_1 band), which is larger than the value found for the Fornax Cluster, which, in fact, has a deeper photometry on the most recent surveys. The distribution of the Sérsic index is shown in Fig. 12(c). The different behaviours for giant and LSB galaxies are evident, pointing to the different morphologies that shape their surface brightness profiles. The histogram has also a unique peak for the whole sample. The mean value for LSB galaxies is $\langle n \rangle = \langle 1/n \rangle = 1.03 \pm 0.45$, which is also larger than the mean Sérsic index found in different environments (de Rijcke et al. 2009; Eigenthaler et al. 2018). The last histogram, Fig. 12(d), shows the distribution of μ_e . There is a main peak at $\mu_e = 25$ mag arcsec $^{-2}$, dominated by LSB galaxies, and a secondary peak at $\mu_e = 23$ mag arcsec $^{-2}$, which seems to lie on the bridge between LSBs and giant galaxies. The distribution of giant galaxies, in turn, has a main concentration at higher surface brightness, at $\mu_e = 21$ mag arcsec $^{-2}$. On the other hand, both redder and bluer galaxies seem to follow the general distribution in effective surface brightness.

4.2 Spatial distribution

The projected spatial distribution of galaxies in the Antlia cluster has been studied by FS90, using photographic plates with a limited magnitude range, but with an interesting dynamical range, which allowed them to define a very precise membership classification based on the galaxies morphology. The elongated distribution of the galaxy population was first suggested by FS90, using the brightest ETGs of the cluster. More recently, X-ray observations reinforced the idea of an axial distribution by mapping the diffuse emission that connects NGC 3258 with NGC 3268 (Pedersen et al. 1997; Nakazawa et al. 2000). Then, Wong et al. (2016) studied the cluster core (a bit larger than the MOSAIC II field 0, see Fig. 5) at 0.5–2 keV and confirmed that the major emission of the X-ray map is centred at NGC 3268 and is elongated on the direction of NGC 3258 (Smith Castelli et al. 2008a; Hess et al. 2015). The latter dominates a younger galaxy subgroup that has recently merged with the first one. Additionally, Hess et al. (2015) report that the Compton-thick Seyfert II NGC 3281 (id 4 on Table 4) has an associated HI absorption region. This galaxy lies – in field 1 – in the direction that connects NGC 3258 with NGC 3268 (see Fig. 6). Now, with a

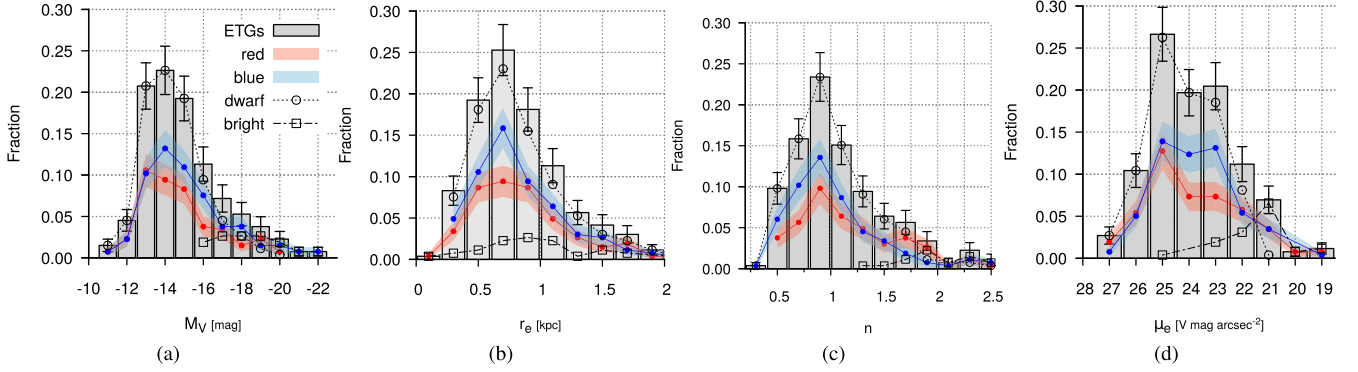


Figure 12. Absolute magnitude and structural parameters distributions (see text).

sample of 265 ETGs covering a magnitude range of $-10 < M_V < -22$ mag, distributed around a cluster-centric distance $R = 700$ kpc, we have built the projected number density profile (Fig. 7). It can be seen that the spatial projected distribution of the ETGs has a main concentration around NGC 3268 and a second concentration around 200 kpc from the cluster centre, close to NGC 3258. Along with this, the angular distribution shows that the ETGs are concentrated along a direction that connects both cD galaxies. On the other hand, redder and bluer ETGs seem to have peculiar characteristics that could be related with the evolutionary state of the cluster, the dynamical relation between the NGC 3258 group and the cluster core, and/or the characteristics of the DM haloes of the cluster itself.

Taking into account the central area of the cluster, up to $R = 300$ kpc, bluer ETGs seem to be more frequent, i.e. on the central field there are more bluer ETGs than redder ones. Within this area, the density profiles of redder and bluer ETGs fall to 0.012 and $0.010 \text{ arcmin}^{-2}$, respectively. While the density of bluer ETGs is strictly decreasing along the whole radial range, redder ETGs display two overdensities near each cD. If this is not a consequence of a detection deficiency, it could be related to the ram-pressure stripping (see e. g. Gunn & Gott 1972) suffered by galaxies within the densest areas of the cluster, which would not happen if they were not linked gravitationally or immersed on the same DM halo (Steinhauser, Schindler & Springel 2016; Zabel et al. 2019). Another interesting feature is that the dwarf-to-giant ratio has a secondary peak near NGC 3258 (at $R \sim 200$ kpc), matching the high-density X-ray emission centred on NGC 3268. Since the Antlia cluster presents a very complex structure, involving a massive cD, associated with a high X-ray emission, and an accreted group of galaxies centred on another cD, it is possible that different physical processes are acting in different degrees. Weinmann et al. (2011) compared the nearest galaxy clusters spatial distributions and colours to semi-analytic models (SAM) on the low-luminosity range of dEs. They collected several samples from the literature of the Virgo, Fornax, Coma, and Perseus clusters and carefully matched them with different SAM samples. The physical properties derived from the SAMs seem to be in good agreement with the observed clusters, even Fornax and Virgo, that have been suggested to be in a non-virialized state. The only parameter that is not well reproduced by the SAMs is the dwarf-to-giant ratio, which gives too high values. They also compared the observed radial density profiles with the simulated average profiles (their fig. 4). The dwarf ETGs on Antlia cover a magnitude range between $-19 < M_V < -11.1$ mag, which implies approximately a number density of $\log(\Sigma) \approx 2.5 \text{ Mpc}^{-2}$ between the central cluster galaxy (NGC 3268) and NGC 3258 (~ 0.2 Mpc), falling down to $\log(\Sigma) \approx 1.7 \text{ Mpc}^{-2}$, around 0.7 Mpc from the centre.

The flattening number of dwarfs around the central area of the cluster, together with the central excess of bluer dwarfs, may be indications of a dynamically young population, as was suggested for the Virgo cluster. This may be produced by stronger tidal disruption suffered by the faintest galaxies in the densest area (which is consistent with the broad velocity distribution, shown in Fig. 9a).

Any correlation between galaxy colours and spatial distribution arises from the colour-density or morphology-density relations (Dressler 1980), which states that there is a correlation between morphological types and the environments in which the galaxies evolve. Many studies confirm that while dEs are commonly found in the centre of clusters, star-forming dwarf irregulars are preferentially located in low-density environments such as cluster outskirts. A more detailed view can be found, for example, in Barazza et al. (2009), who study the population of dEs in the multiple-cluster system Abell 901/902 (at $z = 0.165$), between $-16 > M_r > -26$ mag. They find that the nucleated dwarf elliptical (dEn) galaxies in their sample have different properties compared to the non-nucleated dE: they are more compact, more concentrated to the clusters centres, and they are rounder. These results seem to be largely due to tidal and ram-pressure stripping (Grebel et al. 2003), and even the pre-processing in galaxy groups (like the NGC 3258 subcluster in Antlia) has strongly affected the dwarf population. They also found that the colours of dEs depend on their location within the cluster, in the sense that redder dEs are concentrated to the cluster centre. This behaviour is exactly opposite to our findings in Antlia, although, as seen in Fig. 7(a), the effect is mostly restricted to the innermost bin, containing a low number of galaxies. Another different feature in the Abell 901/902 system is its small colour gradient, as already reported in other galaxy clusters and groups. In Fornax, Venhola et al. (2017) speculated that the galaxies that have spent enough time in the core of the cluster should suffer the removal of their cold gas by ram-pressure stripping. So, this could affect the colour of the central galaxies. The constant colour distribution along the Antlia cluster, until 60 arcmin (~ 600 kpc), may indicate that Antlia is on a different stage on its evolutionary process.

4.3 Radial velocity distribution

The substructures of the Antlia Cluster have also been studied by Hess et al. (2015) in relation to HI and star formation regions. Based on published radial velocities, they measured the relative kinematic deviation from the mean velocity that represents the entire cluster. Besides the obvious structures associated to NGC 3268 (which was stated as cluster centre) and NGC 3258, they proposed three more kinematic structures that coexist in the cluster (see their fig. 7, Hess

Table 5. Antlia UDG candidates.

	$C - T_1$ (mag)	r_e (kpc)	μ_e (mag arcsec $^{-2}$)	n	R (kpc)
Group 1	1.58 ± 0.34	2.03 ± 0.54	24.51 ± 0.27	1.29 ± 0.38	33.35 ± 9.51
Group 2	1.81 ± 0.30	1.81 ± 0.22	26.53 ± 0.42	1.51 ± 0.60	19.85 ± 9.07

et al. 2015), which could be analysed together with our Fig. 9: the substructure (I) (at coordinates 10:29:12.00, $-34:30:00$) lies in Field 2, on the cluster outskirts. Within this area, the most extreme radial velocity galaxies, i.e. FS90 152, FS90 101, are found. The substructure (II) (at 10:30:24.00, $-35:36:00$) is located at the cluster centre and is associated with FS90 226 (id 6) that has a radial velocity representative of the brightest members. The substructure (III) (at 10:30:24.00, $-35:12:00$), located in Field 0, is aligned with the direction connecting NGC 3258 and NGC 3268 and agrees with the position of one of the centroids calculated with the k-means method.

The radial velocity histogram (Fig. 9a) shows that the bluer galaxies that numerically dominate the innermost radial bin have a broad range of velocities (between the membership limits), which may be interpreted as a result of the environment effects produced by the merger process blurring out the substructures in the area of the massive galaxies (see also Hess et al. 2015). A different possible explanation was suggested by Caso & Richtler (2015), using just a sample of radial velocities.

4.4 Luminosity function

The Antlia LF, which was described in Section 3.5 with a Schechter fit: $\alpha = -1.37 \pm 0.03$, and $M^* = -21.70 \pm 1.36$, can be compared with those obtained in Fornax (Venholá et al. 2019) and Virgo clusters (Ferrarese et al. 2016): $\alpha = -1.31 \pm 0.07$ and $\alpha = -1.34^{+0.017}_{-0.016}$, respectively. It is important to note that the Virgo and Fornax values are completeness corrected and obtained from deeper images than ours in the present paper. Other different environment galaxy systems are Hydra I (Misgeld, Mieske & Hilker 2008) and Centaurus (Misgeld, Hilker & Mieske 2009) clusters that have shallower faint-end slopes: -1.13 ± 0.04 and -1.08 ± 0.03 , respectively. These values are in agreement with the LF slope of the Fornax cluster centre, when dEs are considered: -1.09 ± 0.10 (Venholá et al. 2019).

The LF shape is closely related with galaxy mergers that are more frequent in galaxy groups because they can change the luminosity distribution of the galaxies involved (Miles et al. 2004). In this sense, the presence of a dip in the LF of high-velocity dispersion clusters can be related with mergers (Lee et al. 2016), which would not be the case of the Antlia cluster, given its smaller velocity dispersion ($\sigma_v = 500 \text{ km s}^{-1}$; Hess et al. 2015).

4.5 UDG candidates

On the LSB regime, between $-17.0 < M_V < -13.5 \text{ mag}$ (Venholá et al. 2017), we found 12 galaxies that present $r_e > 1.5 \text{ kpc}$ and $\mu_0 > 25.0 \text{ mag arcsec}^{-2}$. These are the characteristics of the so-called UDGs that were found in several clusters and groups (Mihos et al. 2015; Muñoz et al. 2015; van Dokkum et al. 2015; Román & Trujillo 2017). In Fig. 4(e), they can be identified as a group of galaxies (indicated as empty pentagons) extending to higher r_e values; their location has also been highlighted with a green shading to clearly show the position of the Antlia UDGs. These galaxies were also shown in Fig. 4(d), where it can be seen that their μ_e

depart from the general trend, reaching almost $27 \text{ mag arcsec}^{-2}$. We show in Table 5 the mean values of the UDGs structural and photometric parameters, separating them into two groups: (i) Group 1 (G1) is composed by FS90 135, 268, 270, and 329 and (ii) Group 2 (G2) by ANTLJ102936-352445, ANTLJ102843-353933, ANTLJ102856-350435.5, FS90 52, 141, 143, 154, and 293. These two groups clearly differ in mean magnitude: $M_V \sim -15.98 \pm 0.28 \text{ mag}$, and $M_V \sim -14.58 \pm 0.28$, respectively, but they share almost the same mean $r_e \sim 1.8 - 2.0 \text{ kpc}$. We performed a principal component analysis of the UDGs parameters that also confirmed the segregation into two groups.

For comparison, we added in Figs 4(d) and (e), the Fornax sample studied by Venholá et al. (2017, which includes LSBs and UDGs) that spans a range of r_e between 1.6 and 11.3 kpc, largely exceeding the range of the Antlia ones. Larger values of r_e are also found in Virgo (Mihos et al. 2015) and Coma (van Dokkum et al. 2015), which could indicate that the particular Antlia evolutionary state, taking into account the ongoing merger with the NGC 3258 group, could have affected the mechanisms by which the UDGs are produced (i.e. tidal interactions, Mancera Piña et al. 2019; Martin et al. 2019).

The colours of the two groups of UDGs in Antlia also underline their differences. The mean colours are $C - T_1 = 1.58 \pm 0.34 \text{ mag}$ and $1.81 \pm 0.30 \text{ mag}$ for G1 and G2, respectively, making G1 bluer and G2 redder than the CMR fit (see equation 3). Since galaxy colours establish constraints on their formation processes and the Antlia UDGs present colours that are in agreement with the CMR, it is possible that both UDG groups share similar formation mechanisms but have different evolutive histories due to the subsequent interactions within the cluster potential.

From the theoretical side, Di Cintio et al. (2017) studied cosmological simulations of isolated galaxies, which include gas outflow processes due to SNe and massive stars; they were able to reproduce the main observed characteristics of UDGs, based on the idea that the internal processes are more efficient than the environmental ones. Their simulations predict colours $B - V = 0.77 \pm 0.12 \text{ mag}$, consistent with the mean colour of the G1 population, $B - V = 0.72 \text{ mag}$ (we used Buzzoni 2005 for the $C - T_1$ to $B - V$ colour transformation), and the mean colour of the UDG sample in Fornax (Venholá et al. 2017). Additionally, the rest of structural parameters (r_e , μ_e) obtained from the simulations are consistent with the UDGs in Antlia with the exception of the Sérsic index.

The G1 galaxies that have the brighter μ_e and bluer colours are distributed on average around $R = 33.35 \pm 9.51 \text{ arcmin}$, which is significantly larger than that for G2 ($R = 19.85 \pm 9.07 \text{ arcmin}$) that coincides with the distance at which NGC 3258 is located.

Finally, G1 galaxies have bluer colours than the CMR, are isolated, and mainly scattered across the cluster area. These UDGs have similar colours than the simulated galaxies in Di Cintio et al. (2017). On the other hand, the G2 is located in the internal part of the cluster, with mean colours redder than the CMR. One possible explanation is that the natural UDG population of each dominant galaxy have suffered different exposition to the environmental processes, due to the merger between the group centred in NGC 3258 and the original population.

5 CONCLUSIONS

In this third paper, which is part of a series (Paper I; Paper II) about the Antlia ETG populations, we have added the surface photometry of ~ 130 galaxies in the magnitude range $-21 < M_V < -11$ mag that have not been previously identified. The majority of the new galaxies occupy the LSB regime; however, some brighter ETGs (mainly S0) confirm the change in slope of the CMR at $M_{T1} \sim -20$ mag (Fig. 3). At the same time, the dispersion of the CMR is not significantly increased with respect to the previous papers. The colour dispersion at each magnitude from $T_1 = 15$ to 19 mag is $\sigma_{(C-T_1)} = 0.16\text{--}0.17$. The dispersion of the Virgo CMR, in the same luminosity regime, is $\sigma_{(g'-z')} = 0.144$ mag (Hamraz et al. 2019), which results to be larger than the dispersion obtained for Fornax and Coma. This larger scatter in the Antlia dEs colours and the most bluer galaxies that we found could indicate that the forming star processes in the LSB population have not finished yet.

We reanalyse the relations between the structural parameters for the Antlia ETGs and, to complement the LSB regime, compare them with those of the Fornax cluster. A linear relation is present between μ_e and M_V (Fig. 4d) that extends from dEs and dSphs up to bright Es and S0, while UDGs and cEs are outliers. The two dominant cD galaxies also fall below the linear fit.

The projected spatial distribution (Figs 5 and 6) shows an elongated distribution of the ETG population in the direction connecting NGC 3258 and 3268. The smoothed X-ray map (Wong et al. 2016) in Fig. 5 has a perfect match with some galaxies in the central region. If we consider the whole MOSAIC images including all galaxies of the final sample, i.e. almost 2.6 deg^2 , the spatial distribution maintains its elongated structure (Fig. 6).

In order to analyse quantitatively the projected distribution, we split the sample into galaxies redder and bluer with respect to the CMR regression, and into dwarf and bright galaxies with respect to $\mu_0 = 19 \text{ mag arcsec}^{-2}$ (V band), counting all galaxies along their distance from NGC 3268, which is adopted as the cluster centre:

- (i) Bluer galaxies are more numerous until the first 200 kpc (Fig. 7a), which corresponds to the position of NGC 3258. Further out, the projected densities of bluer and redder galaxies are quite similar. Moreover, the azimuthal distribution (Fig. 7b) reaches a maximum density in the direction towards (and opposite to) NGC 3258, confirming the elongated structure mentioned above.
- (ii) The redder galaxies seem to be more concentrated to the bright galaxies in the cluster, and their radial velocities clearly occupy the centre of the histogram between $V_R = 2000$ and 3000 km s^{-1} .
- (iii) The bright galaxies tend to be concentrated towards the cluster centre, while the azimuthal distribution remains homogeneous. The case is different for the dwarf galaxies, which trace the general trend of the ETG population.

There does not appear to be any colour gradient along the cluster-centric radius (Fig. 8), being the mean colour $\langle (C - T_1)_0 \rangle \sim 1.69 \pm 0.09$ mag up to 700 kpc. However, the mean r_e of the galaxies decreases from the centre outwards, with a stronger slope outside of the position of NGC 3258.

The Antlia LF (Fig. 10) can be well fitted with a single Schechter function, with a slope comparable with those of the Fornax and Virgo clusters. Additionally, the low number of bright ETGs, the weak X-ray emission in the cluster centre, the ongoing merger of NGC 3258 group, and the apparently early state of virialization may be the main causes of the Antlia LF shape and the lack of a dip at intermediate magnitudes.

We found 12 UDG candidates that meet the criteria described by van Dokkum et al. (2015). All UDG candidates have $r_e < 3$ kpc, which is not the general case as larger values have been found in other clusters. UDGs in Antlia can be grouped in two different families, characterized by different mean absolute magnitude and effective surface brightness. These two groups also present different colours and projected spatial distributions, possibly indicating a distinctive origin or evolution mechanism, due to the ongoing merger with the group centred on NGC 3258.

ACKNOWLEDGEMENTS

We thank to an anonymous referee for constructive remarks. This work was funded with grants from Consejo Nacional de Investigaciones Científicas y Técnicas de la República Argentina, Agencia Nacional de Promoción Científica y Tecnológica, and Universidad Nacional de La Plata, Argentina. LPB, JPCaso, and MG: Visiting astronomers, Cerro Tololo Inter-American Observatory, National Optical Astronomy Observatories, which are operated by the Association of Universities for Research in Astronomy, under contract with the National Science Foundation.

DATA AVAILABILITY

The data underlying this article are available in the article and in its online supplementary material.

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SUPPORTING INFORMATION

Supplementary data are available at *MNRAS* online.

Table 3. Catalogue of galaxies in the Antlia cluster.

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